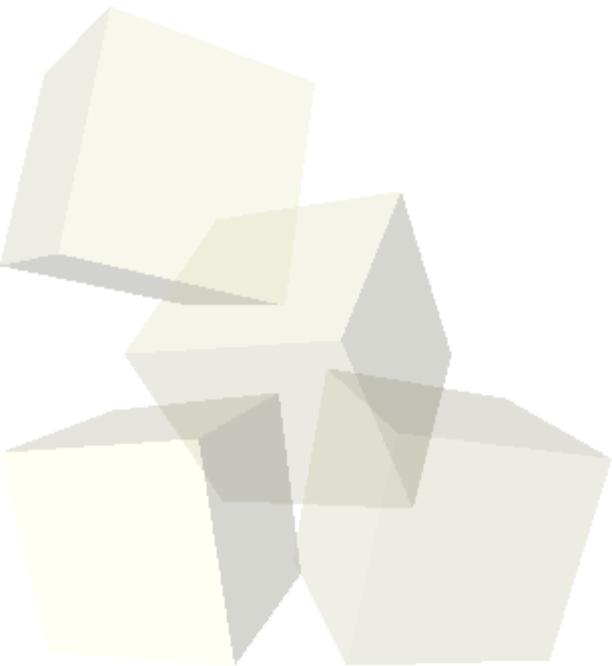




15th Capra Ranch Meeting

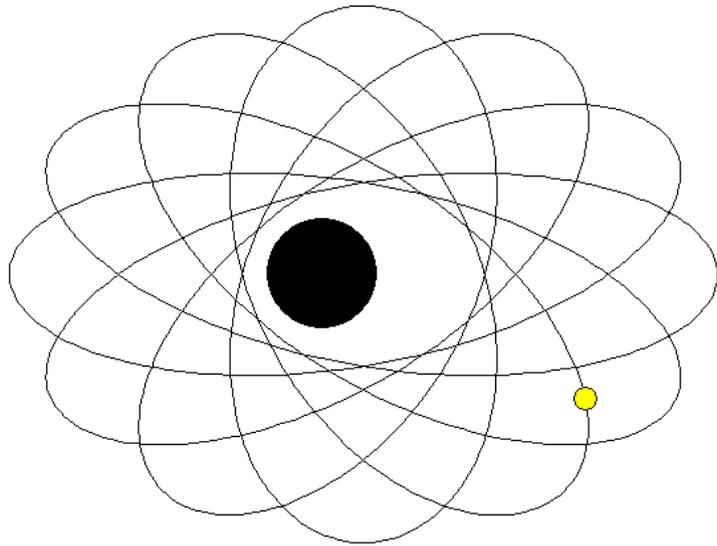
Self-force driven inspiral of a scalar point particle into a Schwarzschild black hole: a progress report

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Extreme Mass Ratio Inspirals



- Solar-mass, compact object moving around a supermassive black hole
 - ◆ energy lost due to radiation leads to orbital decay
 - ◆ possible source for space based GW antenna
- 10^5 wave cycles / year
 - ◆ below noise level
 - ◆ accurate modeling to detect signal
 - ◆ encode geometry around central object
- employ **small mass ratio approach**
 - ◆ first order expansion in μ/M
 - ◆ no weak-field or slow motion assumption

Scalar self force in Schwarzschild

field sourced
by particle

$$\square\Phi = -4\pi\rho$$

$$m(\tau)\ddot{x}^\alpha = q(g^{\alpha\beta} + u^\alpha u^\beta)\nabla_\beta\Phi$$

$$\dot{m} = -qu^\alpha\nabla_\alpha\Phi$$

loss of rest mass
to monopole radiation

particle moves under
influence of field

- toy model for gravity
 - ♦ classical, massless field
 - ♦ “particle” is stellar sized object in orbit around black hole
- self-consistent evolution must evolve field and particle simultaneously



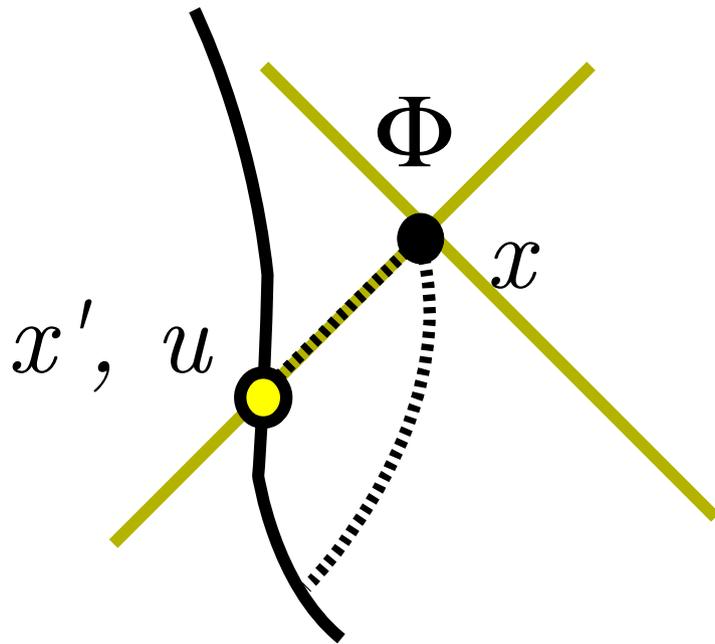
■ Motivation

- ♦ current mode-sum approach calculates the self-force on a geodesic
- ♦ one inspiral calculation requires data from a large number of geodesics \Rightarrow is calculated as a postprocessing step
- ♦ inefficient, since a large bank of self-force templates has to be computed in advance

■ Idea

- ♦ calculate the self-force self-consistently along an accelerated world line and use it to evolve forward in time
- ♦ calculation of self-force and evolution of orbit occur at the same time

Detweiler-Whiting decomposition



- waves travel both
 - ♦ **directly** along the light cone
 - ♦ by scattering off curvature in the **tail**
- particle interacts with its own radiation

- direct piece is singular and must be removed
- tail piece is regular and solely responsible for the force

controls motion $\rightarrow \nabla_{\alpha} \Phi^R = \nabla_{\alpha} \Phi - \nabla_{\alpha} \Phi^S$

compute numerically $\rightarrow \nabla_{\alpha} \Phi$

known analytically $\rightarrow \nabla_{\alpha} \Phi^S$

■ Mode sum regularization

- ♦ regularizes after computing Φ
- ♦ efficient in Schwarzschild, hard to extend to Kerr
- ♦ first method to be used successfully [Barack, Mino, Nakano, Ori, Sasaki 2002]

■ Effective source method

- ♦ regularizes source term in wave equation
- ♦ full 3D simulation for field, extension to Kerr is simpler
- ♦ first method to compute self-consistent motion [Diener, Vega, Wardell, Detweiler 2011]

■ m-mode regularization

- ♦ combines aspects of effective source and mode sum regularization
- ♦ designed to work in Kerr [Barack, Golbourn, Sago 2007]

■ Green function methods



Mode sum regularization

- Spacetime is spherically symmetric
 - ♦ decompose field into spherical harmonic modes
 - ♦ modes decouple
 - ♦ each mode is finite at location of particle

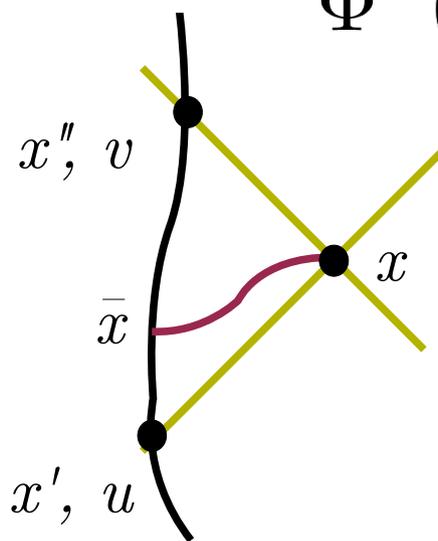
$$\nabla_{\alpha} \Phi^S = \sum_{\ell} (\nabla_{\alpha} \Phi^S)_{\ell} P_{\ell}$$

$$(\nabla_{\alpha} \Phi^S)_{\ell} = (\ell + 1/2)A_{\alpha} + B_{\alpha} + \frac{C_{\alpha}}{\ell + 1/2} + \frac{D_{\alpha}}{(2\ell - 1)(2\ell - 3)} + O(\ell^{-3})$$

Annotations:
- "vanishes" with an arrow pointing to the C_{α} term.
- "independent of ℓ " with arrows pointing to the A_{α} and B_{α} terms.

(numerical $\nabla_{\alpha} \Phi$ must match this structure)

Singular field near the world line



The diagram shows a curved world line. A point on the world line is labeled x', u . A point on the world line is labeled x'', v . A point on the world line is labeled \bar{x} . A point on the world line is labeled x . A red curve connects \bar{x} and x . A yellow line connects x', u and x . A yellow line connects x'', v and x .

$$\Phi^S(x) = \frac{q}{2r} U(x, x') + \frac{q}{2r_{\text{adv}}} U(x, x'') - \frac{1}{2} q \int_u^v V(x, z) d\tau$$

x', x'' retarded/advanced point

- Expand (bi-)tensors in terms of $\sigma^{\bar{\alpha}} \equiv \nabla^{\bar{\alpha}} \sigma(x, \bar{x})$ (covariant)

$$A(x, \bar{x}) = A(\bar{x}) + A_{\bar{\alpha}}(\bar{x}) \sigma^{\bar{\alpha}} + \dots$$

- Expand σ^{α} in terms of $(x - \bar{x})^{\bar{\alpha}}$ (not covariant)

$$\sigma^{\bar{\alpha}} = (x - \bar{x})^{\bar{\alpha}} + B_{\bar{\beta}\bar{\gamma}}^{\bar{\alpha}} (x - \bar{x})^{\bar{\beta}} (x - \bar{x})^{\bar{\gamma}} + \dots$$



Changes due to acceleration

- Covariant local expansion of the singular field
 - Synge's world function σ links points on the world line
 \Rightarrow acceleration and higher derivatives appear in its expansion along the world line
 - Retarded and advanced times depend on $\sigma \Rightarrow$ acceleration appears
- Coordinate expansion of bitensors
 - Unchanged as the point x on the world line is arbitrary

$$A_{(\mu)} = \hat{A}_{(\mu)}(x^\alpha, u^\alpha) \text{sign}(\Delta)$$

$$B_{(\mu)} = \hat{B}_{(\mu)}(x^\alpha, u^\alpha, a^\alpha)$$

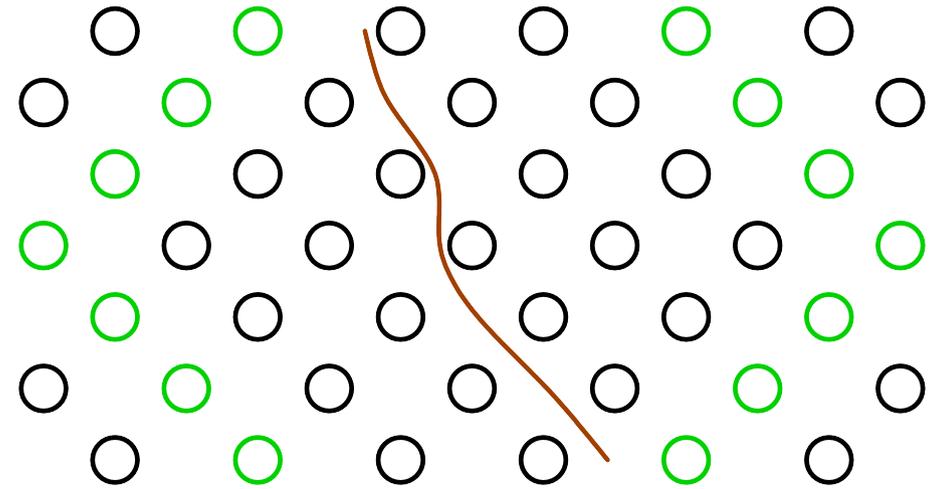
$$C_{(\mu)} = 0$$

$$D_{(\mu)} = \hat{D}_{(\mu)}(x^\alpha, u^\alpha, a^\alpha, \dot{a}^\alpha)$$



$$\square_{\text{flat}}^{1+1} (r\Phi_{\ell m}) - V \Phi_{\ell m} = S_{\ell m} \delta(r - r(\tau))$$

- Fourth-order accurate algorithm
- No boundary conditions are enforced, instead the grid matches the domain of dependence in each timestep
- No physical initial data is specified, we wait until the initial radiation contents has propagated away
- Both the field and the source term are evolved concurrently using a predictor-corrector scheme





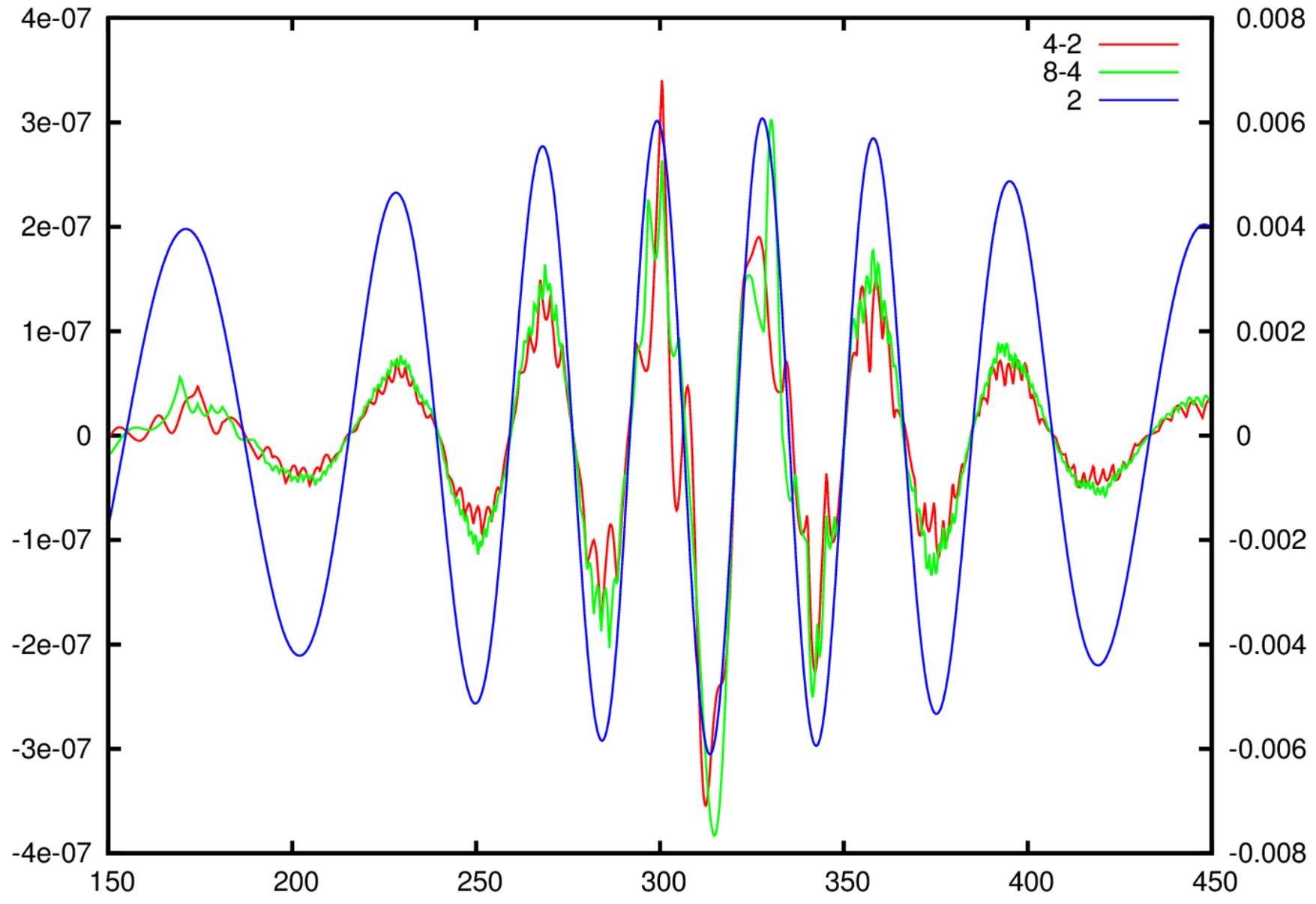
Changes to handle acceleration

- field evolved using the original fourth-order accurate algorithm of [Lousto 2009]
 - ◆ time update requires only two time slices
 - ◆ straightforward to adapt to accelerated motion
 - ◆ fast
- particle evolved in step with field using an Adams-Bashforth-Moulton multistep timestepper
 - ◆ only uses time steps for which field values are available
 - ◆ 4th order accurate in time
- extraction of field value at particle location uses partial information on jumps. Only jumps independent of acceleration are used



Code convergence

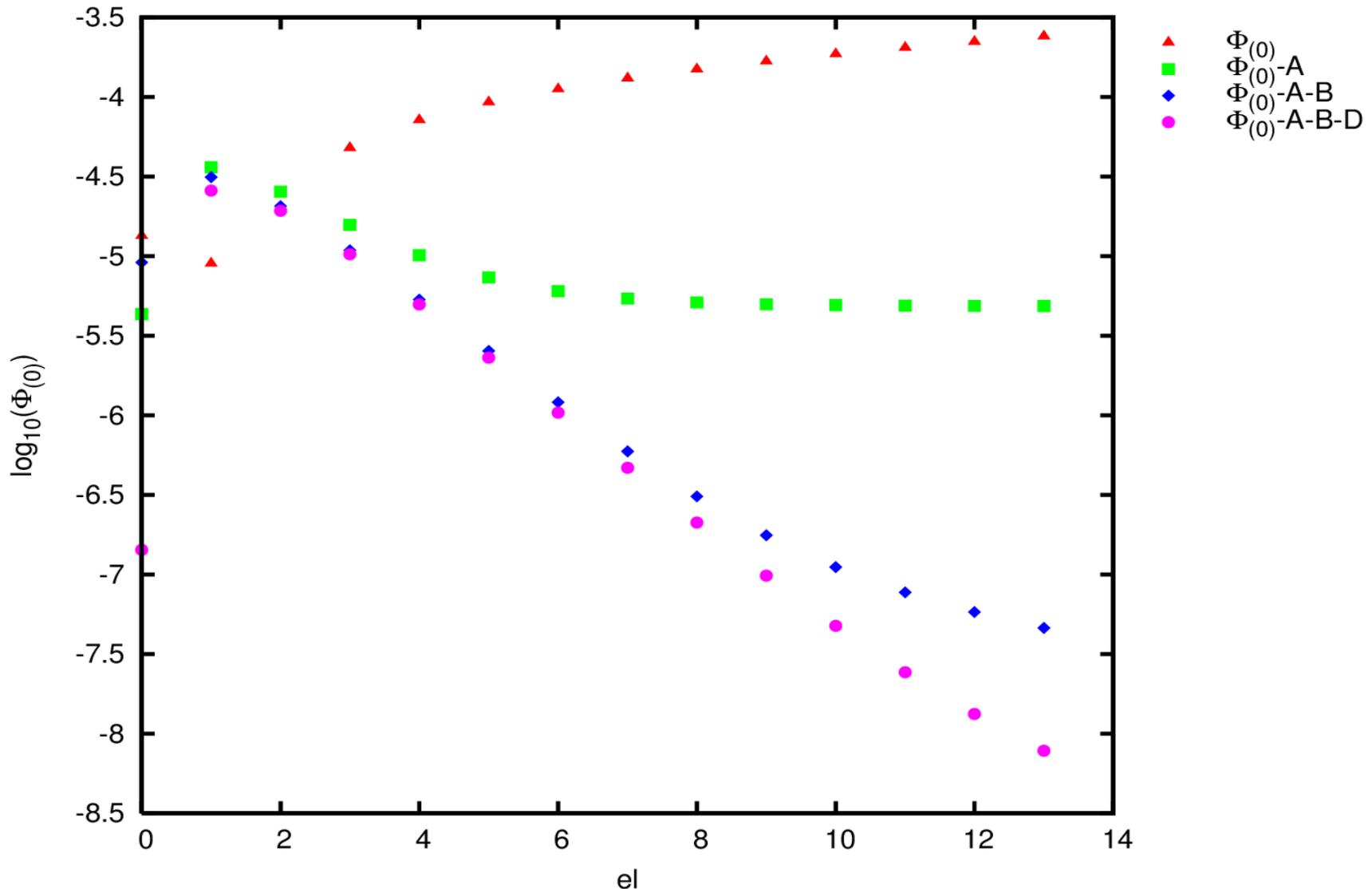
Regression of order 4
q=0.1 el=5 em=3



Mode falloff for geodesic orbits

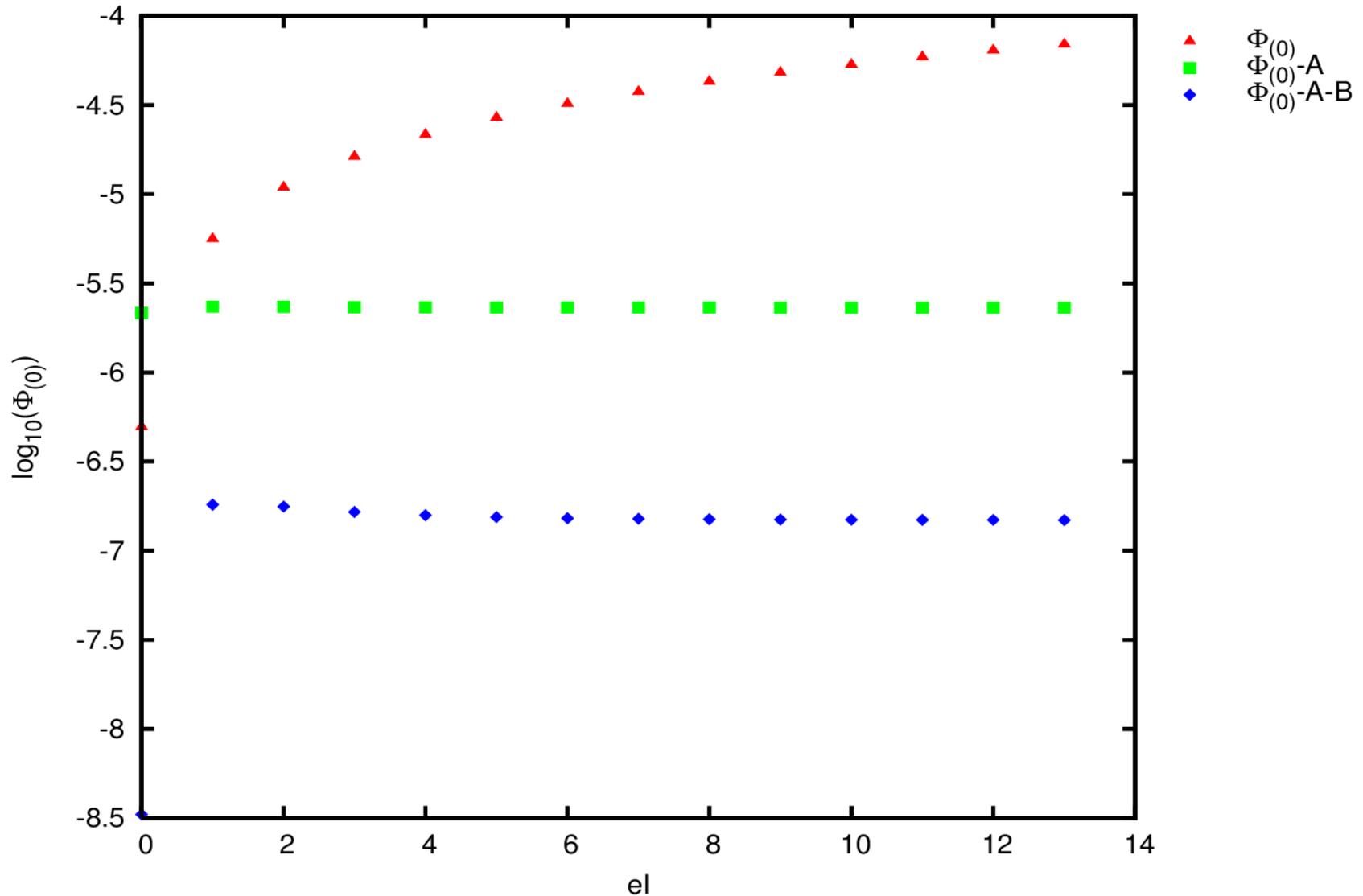
$$(\nabla_{\alpha}\Phi^S)_{\ell} = (\ell + 1/2)A_{\alpha} + B_{\alpha} + \frac{C_{\alpha}}{\ell + 1/2} + \frac{D_{\alpha}}{(2\ell - 1)(2\ell - 3)}$$

l.622713, e = 0.500000, p = 7.200000, q = 0.050000, Delta_t = 0.00416667, Delta_rstar = 0.00833333, t_extract = 10



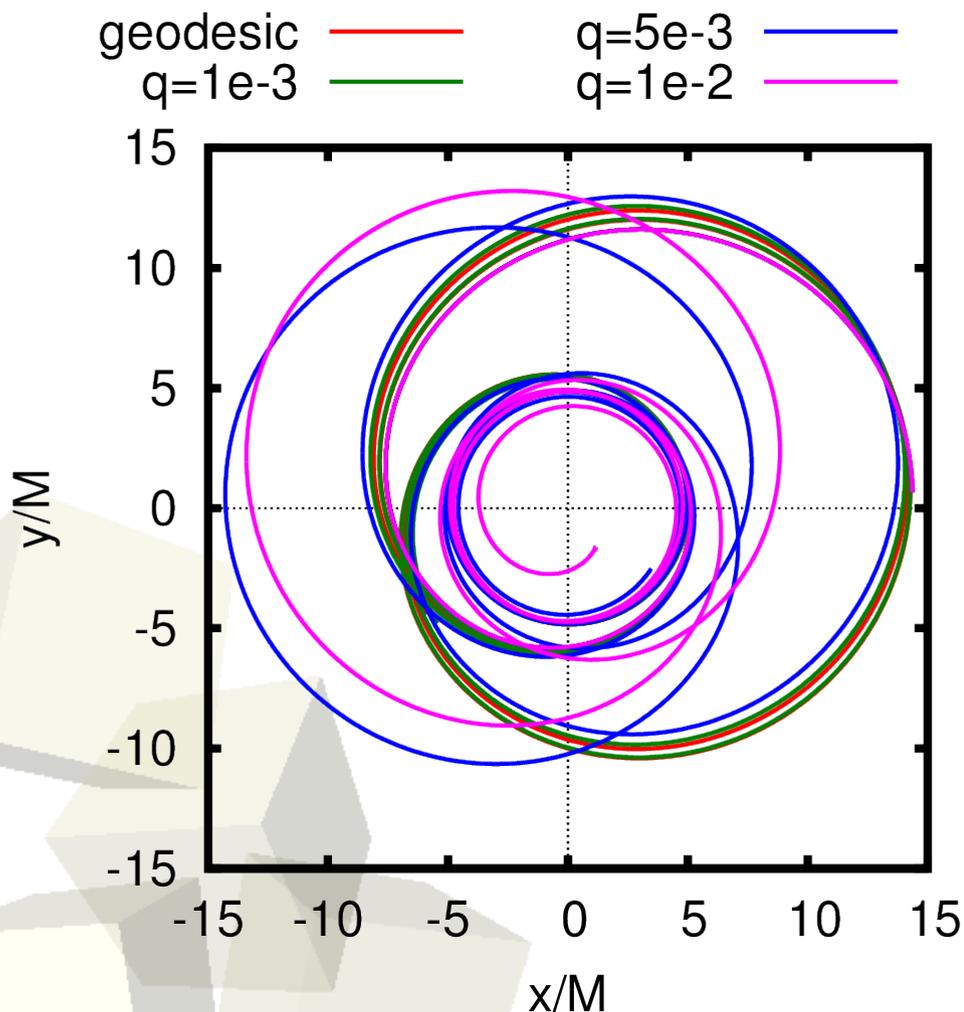
$$(\nabla_{\alpha} \Phi^S)_{\ell} = (\ell + 1/2)A_{\alpha} + B_{\alpha} + \frac{C_{\alpha}}{\ell + 1/2} + \frac{D_{\alpha}}{(2\ell - 1)(2\ell - 3)}$$

= 3.622713, e = 0.500000, p = 7.200000, q = 0.005000, Delta_t = 0.00416667, Delta_rstar = 0.00833333, t_extract =





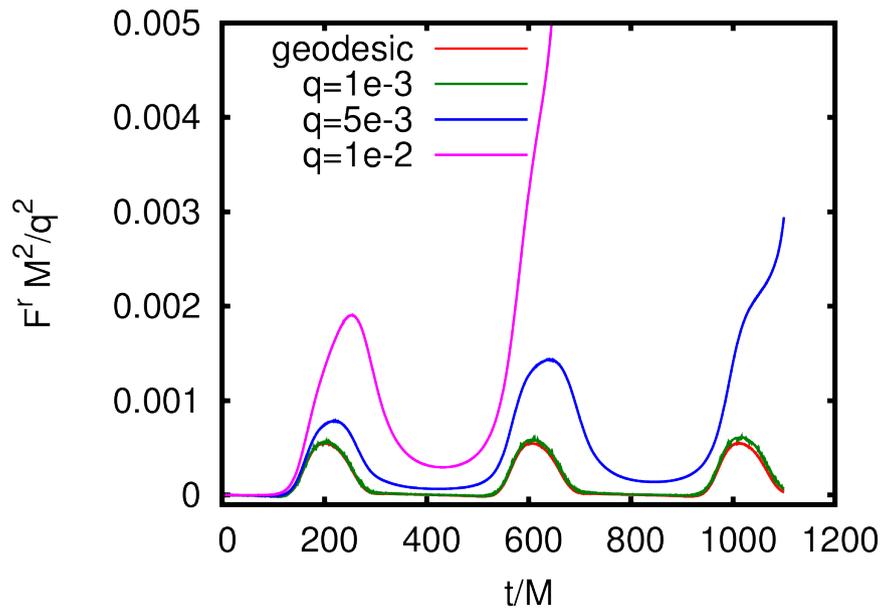
zoom-whirl
 $p = 7.2$ $e = 0.5$



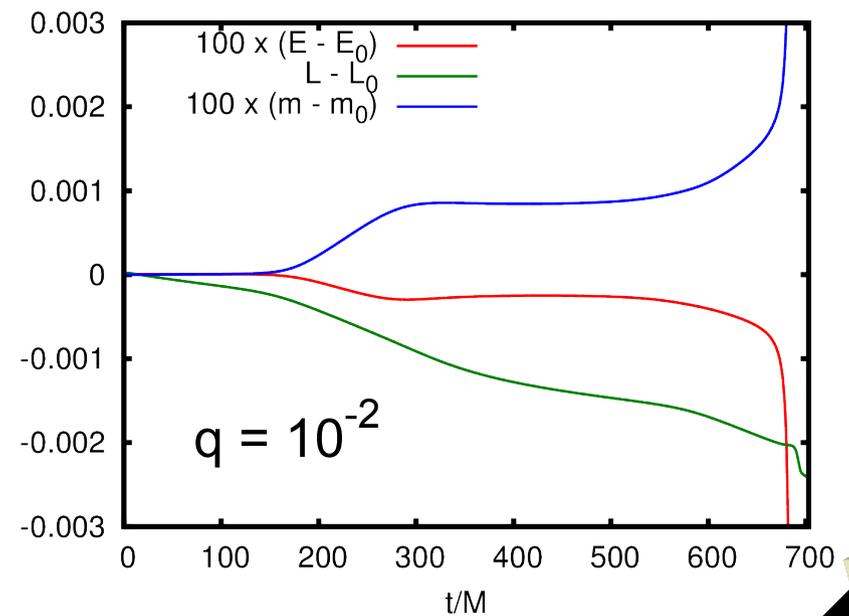
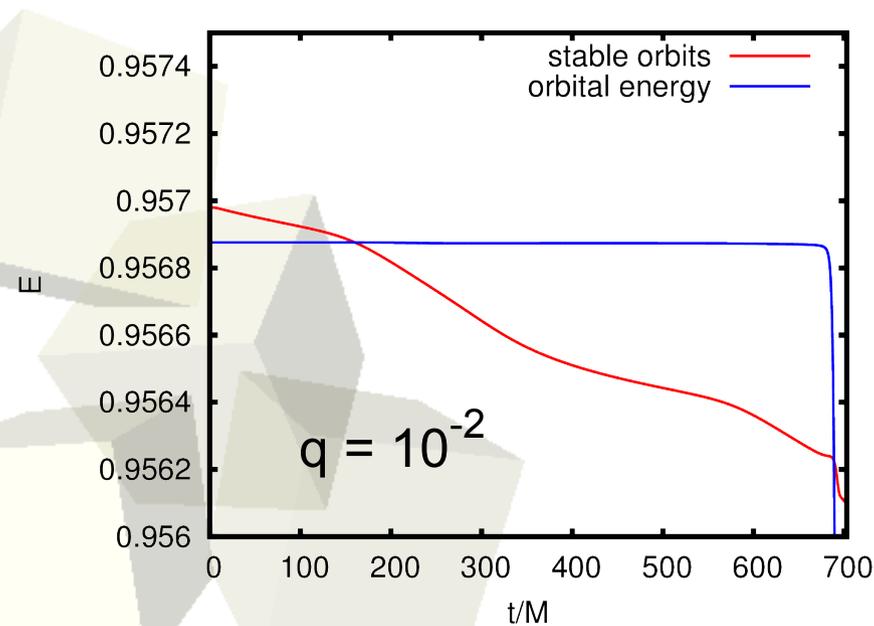
- zoom whirl orbit
 - ◆ copious radiation emitted during whirl phase
 - ◆ penetrates deep within the strong field region
- self force computed locally
- increased perihelion advance due to self force
- sudden transition from inspiral to plunge

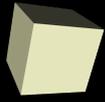


Variable constants of motion



- charge controls number of orbits before plunge
- mass loss due to scalar radiation
- E, L oscillate





- self-consistent evolution of charge under the influence seems possible but there are still bugs in the code
- self-force likely increases perihelion advance
- onset of plunge once
 - ◆ angular momentum is sufficiently low
 - ◆ particle on inbound trajectory

